

RECEIVED: February 6, 2025

REVISED: May 23, 2025

ACCEPTED: June 17, 2025

PUBLISHED: July 22, 2025

Study of tau neutrinos and non-unitary neutrino mixing with the first six detection units of KM3NeT/ORCA



The KM3NeT collaboration

E-mail: km3net-pc@km3net.de, cerisy@cppm.in2p3.fr,
nicole.geisselbrecht@fau.de, lastoria@lpccaen.in2p3.fr

ABSTRACT: Oscillations of atmospheric muon and electron neutrinos produce tau neutrinos with energies in the GeV range, which can be observed by the ORCA detector of the KM3NeT neutrino telescope in the Mediterranean Sea. First measurements with ORCA6, an early subarray corresponding to about 5% of the final detector, are presented. A sample of 5828 neutrino candidates has been selected from the analysed exposure of 433 kton-years. The ν_τ normalisation, defined as the ratio between the number of observed and expected tau neutrino events, is measured to be $S_\tau = 0.48^{+0.5}_{-0.33}$. This translates into a ν_τ charged-current cross section measurement of $\sigma_\tau^{\text{meas}} = (2.5^{+2.6}_{-1.8}) \times 10^{-38} \text{ cm}^2 \text{ nucleon}^{-1}$ at the median ν_τ energy of 20.3 GeV. The result is consistent with the measurements of other experiments. In addition, the current limit on the non-unitarity parameter affecting the τ -row of the neutrino mixing matrix was improved, with $\alpha_{33} > 0.95$ at the 95% confidence level.

KEYWORDS: Beyond Standard Model, Neutrino Detectors and Telescopes (experiments), Oscillation

ARXIV EPRINT: [2502.01443](https://arxiv.org/abs/2502.01443)

Contents

1	Introduction	1
2	The KM3NeT/ORCA detector	3
2.1	Data taking, event reconstruction, and trigger algorithms	3
2.2	Data processing and quality selection	4
2.3	Monte Carlo simulations	4
3	Physics context	5
3.1	From the ν_τ normalisation to the ν_τ CC cross section measurement	5
3.2	Non-unitary neutrino mixing	6
4	Event selection	9
5	Analysis method	9
6	Results	12
6.1	ν_τ CC cross section measurement	12
6.2	Probing unitarity	16
7	Comparison with other experiments and conclusions	18
	The KM3NeT collaboration	25

1 Introduction

The experimental observation of neutrino oscillations in 1998 [1] provided evidence that neutrinos have non-zero mass. Neutrino flavour eigenstates (ν_e, ν_μ, ν_τ) are linear combinations of the mass eigenstates (ν_1, ν_2, ν_3) expressed through the Pontecorvo-Maki-Nakagawa-Sakata (PMNS) mixing matrix [2, 3]. The oscillation probabilities are a function of the distance travelled by neutrinos, their energy, the mixing matrix and the mass-squared differences between the mass states, $\Delta m_{ij}^2 = m_i^2 - m_j^2$. However, the Standard Model of elementary particles does not provide a mechanism describing the origin of such masses, and neutrino physics is a first field showing a clear necessity for physics beyond the Standard Model.

From the experimental data gathered so far, a coherent picture of the neutrino mixing has been achieved. The PMNS matrix (U_{PMNS}) is defined by three mixing angles ($\theta_{12}, \theta_{13}, \theta_{23}$), a charge-parity (CP) violation phase (δ_{CP}), and is assumed to be unitary. The values of the oscillation parameters are currently known with percentage-level precision [4–6]. Critical questions regarding the oscillation mechanism still need to be addressed. Providing those answers, with the determination of the neutrino mass ordering by clarifying the sign of Δm_{31}^2 and the precise measurement of the δ_{CP} -phase to investigate the presence of CP violation in the leptonic sector, would open a way to explore new physics. While the quark mixing matrix has been extensively tested with robust precision, the unitarity of the PMNS matrix,

assumed in many oscillation analyses [7, 8], still has to be experimentally proven. Under the PMNS matrix unitarity hypothesis, solar [9, 10], reactor [11–14], and atmospheric [1, 15, 16] neutrino experiments mostly exploited the ν_e , $\bar{\nu}_e$, and ν_μ , $\bar{\nu}_\mu$ disappearance channels, whereas the high-statistics in ν_e and $\bar{\nu}_e$ appearance channels have been investigated in long-baseline accelerator experiments [7, 8].

The detection and study of τ neutrinos¹ is challenging. Their existence was suggested in 1977 [17], while their first experimental observation was reported only in 2000 [18]. The main detection channel is the τ lepton produced via charged-current (CC) interactions, with a production energy threshold of about 3.5 GeV and a suppressed cross section near this threshold. At accelerators, ν_τ are created through the decays of heavy charmed hadrons. In the atmosphere, ν_τ can be observed mainly via the oscillation channel, largely taking place above the production threshold, which becomes accessible for a baseline longer than ~ 2000 km. However, the short lifetime of the τ -lepton requires excellent particle identification techniques for a clean ν_τ event reconstruction [18, 19]. Overall, the main current limitation is the statistics. Therefore, in contrast with the current precision on the PMNS matrix elements in the electron and muon rows, the τ -row, remains the least constrained. This difference is enhanced when allowing for non-unitarity [20–22]. In the standard three neutrino framework (3ν -paradigm), 12 conditions on the PMNS matrix elements exist; six of them are given by the sum of the elements squared in each row (referred to as *normalisation* and expected to be equal to 1). A plethora of beyond the Standard Model extensions are proposed to validate the 3ν -paradigm. Moreover, the tau neutrino interaction model and cross section currently lack precise experimental data, in order to validate the theoretical prediction [23]. In this context, the ν_τ appearance, as direct detection of ν_τ from the oscillation mechanism, is a golden channel to constrain the ν_τ CC cross section.

Several running and next-generation neutrino experiments are exploiting different techniques to reach a deeper understanding of tau neutrinos [24] from the GeV to the EeV energy scale. This paper investigates the properties of atmospheric tau neutrinos, using data collected with a partial configuration of the KM3NeT/ORCA neutrino telescope (5% of the nominal volume, in the following referred to as ORCA6), from January 2020 to November 2021. Currently under construction in the Mediterranean Sea offshore the south of France, KM3NeT/ORCA is optimised for the study atmospheric neutrino oscillations in the 1–100 GeV energy range. Considering the foreseen final instrumented volume, the strong matter effect, the accessible low-energy detection threshold, and the good angular and energy resolutions, the $\nu_\mu \rightarrow \nu_e$ appearance channel will allow for probing the neutrino mass ordering by studying the resonant oscillation in the Earth’s mantle [25]. The ν_μ disappearance is the dominant channel and contributes to a precise measurement of the atmospheric oscillation parameters Δm_{31}^2 and θ_{23} . Additionally, KM3NeT/ORCA will have good sensitivity to the ν_τ appearance channel thanks to its large accessible statistics, ~ 3000 ν_τ CC/year in the final detector volume [25, 26].

The structure of the paper is organised as follows. The KM3NeT/ORCA telescope and the used data sample are described in section 2; Monte Carlo simulations and the data taking conditions are described in section 2.1 and 2.3. In section 3, the theoretical aspects of the

¹In this paper, the term *neutrino* generally refers to both neutrinos and antineutrinos if not specified otherwise.

ν_τ appearance channel investigation are explained, with a focus on the impact on the ν_τ cross section measurement in section 3.1 and on the theory behind the non-unitarity test of the neutrino mixing matrix in section 3.2. The event selection and analysis method are reported in section 4 and section 5, respectively. The results are summarised in section 6: the cross section measurement and results on non-unitarity constraints are reported in section 6.1 and 6.2, respectively. A conclusive summary describing the physics contribution of ORCA6 analysis to already existing results is reported in section 7.

2 The KM3NeT/ORCA detector

The KM3NeT Collaboration is building two large-volume water Cherenkov neutrino detectors in the Mediterranean Sea: KM3NeT/ARCA (Astroparticle Research with Cosmics in the Abyss) and KM3NeT/ORCA (Oscillation Research with Cosmics in the Abyss). The two detectors have complementary physics programs and explore neutrinos in a wide energy range, from MeV to PeV [26]. KM3NeT/ARCA, under construction 100 km offshore Portopalo di Capo Passero (Sicily, Italy), at a depth of 3500 m, is optimised for high energy astrophysical neutrino detection in the TeV-PeV energy range. KM3NeT/ORCA is being built at 2450 m depth, 40 km offshore Toulon (France). Both detectors consist of a 3-dimensional array of photomultiplier tubes (PMTs) collecting the Cherenkov radiation induced in water by relativistic charged particles emerging from neutrino interactions with the seawater and seabed. The detector’s key component is a spherical structure, the Digital Optical Module (DOM), hosting 31 PMTs, ensuring 4π coverage for the detection of Cherenkov photons [27]. A set of 18 DOMs are integrated into a Detection Unit (DU), kept vertically from the seabed thanks to buoys completing the top of each DU and the buoyancy of the DOMs themselves. The horizontal and vertical distances between optical sensors in the two neutrino telescopes are optimised to facilitate their main physics goals. In order to detect GeV neutrinos with KM3NeT/ORCA, its vertical inter-DOM and horizontal inter-DU distances are ~ 9 m and ~ 20 m, respectively. The number of installed DUs determines the fiducial detector mass and will reach about 7 Mton (115 DUs) in the final design geometry. Thanks to its modular structure, KM3NeT/ORCA can detect neutrinos with a partially instrumented volume, which at the time of writing, covers about 20% of the nominal volume.

2.1 Data taking, event reconstruction, and trigger algorithms

Atmospheric neutrinos are the main signal in KM3NeT/ORCA. They are the decay products of kaons and pions created in the collision between cosmic rays and nuclei in the Earth’s atmosphere. Contaminating the neutrino signature, two main background sources are present: atmospheric muons and environmental optical noise. The former consists of downgoing muons from extensive air showers. The latter is mostly caused by the radioactive decay of ^{40}K in seawater and bioluminescence present in the detector surroundings.

The KM3NeT detector readout system [28] is based on the *all-data-to-shore* concept, so that all analogue signals from the PMTs exceeding a tunable threshold (typically 0.3 photoelectrons) are digitised and sent to shore for real-time processing. The data stream is processed using different levels of trigger conditions, including time and geometrical causality of the triggered hits. The stored optical data (the “*event*”) contains the time and the

time-over-threshold of each analogue pulse (jointly referred to as “*hit*”) in a time window of a few μs around those hits that satisfy one of the trigger conditions. This procedure provides a 10^5 reduction factor to the data stream before storing it on disk.

The event reconstruction is based on maximum-likelihood algorithms to extracting the relevant information of each event based on the space-time distribution of PMT signals in the detector. Depending on the neutrino flavour and the type of interactions (CC and neutral-current, NC), two main event topologies can be identified: *track*-like and *shower*-like. The first topology occurs every time a muon is produced as a secondary particle; the latter, occurs in all the other cases. The maximum likelihood algorithms assume, for each event, either a track or a shower hypothesis and provide an estimation of the direction and energy [29]. In ORCA6, 70% of track-like events are correctly reconstructed for events with energy above $\sim 30\text{ GeV}$; below this energy, events are identified as shower-like.

2.2 Data processing and quality selection

The data stream is divided into *runs*, with a typical duration of about 6 hours. The run duration is chosen to facilitate file transfer to KM3NeT Tier-1 computing centres and subsequent analysis. Summary data containing the rates of all PMTs in the detector are stored with a sampling frequency of 10 Hz. This information is used in the simulation and the reconstruction of the events on a run basis.

For the ORCA6 dataset, additional data quality criteria are applied to exclude runs acquired in particularly unstable conditions, due to periods of high bioluminescence, timing accuracy issues, and high trigger rates corresponding to 4.2% of the raw data sample (more details on the quality criteria applied for the run selection are given in [29]). The results reported in this paper are obtained from a data sample including the reconstruction of both track-like and shower-like topologies, which corresponds to an exposure of 433 kton-years (510 days of detector uptime). Starting from this dataset, events are selected and classified to reject the main background and identify the neutrino sample as detailed in section 4.

2.3 Monte Carlo simulations

Detailed Monte Carlo (MC) simulations of atmospheric neutrinos and atmospheric muons are used to evaluate the KM3NeT/ORCA detector response. A *run-by-run* approach [30] is used to mitigate the time-dependent impact of data taking. This approach exploits information on the active/inactive PMTs, the run setup, and the level of environmental optical background. The simulation chain [29] uses the GENIE-based [31] software package gSeaGen [32] for neutrino-induced interactions in seawater with the tune G18_02a_00_000. Atmospheric neutrinos in the 1 GeV–10 TeV energy range are simulated; all neutrino flavours and both CC and NC interactions are considered. MC events are weighted according to the Honda model [33], using the atmospheric neutrino flux estimated in the Northern hemisphere (Fréjus site). The path of neutrinos through the Earth is simulated using a density profile model based on the Preliminary Reference Earth Model (PREM) [34], consisting of 15 layers of different densities. The neutrino oscillation probabilities are calculated using the OscProb software [35] where also the non-unitarity neutrino mixing model (see section 3.2) has been implemented. As detailed in [26], secondary particles produced from neutrino interactions with

seawater are propagated using GEANT4-based custom KM3NeT software packages; similarly, the light propagation is simulated by taking into account the absorption and scattering lengths as well as the PMT quantum efficiency. Depending on the neutrino energy, the light propagation is either fully simulated or obtained through look-up tables to reduce CPU time consumption. The atmospheric muon generation is based on the MUPAGE package [36], while their propagation in seawater uses a custom software package [37].

3 Physics context

Due to the very short τ lepton lifetime, attempting an event-by-event ν_τ identification in KM3NeT/ORCA is currently not possible. Nevertheless, given the atmospheric neutrino flux composition [38], the current active volume, and the GeV-scale sensitivity, the tau neutrino appearance can be quantified on a statistical basis. The ν_τ normalisation [19, 23, 39] is defined as the ratio between the observed and expected number of tau neutrinos. When accounting for the charged-current ν_τ events only, the ν_τ normalisation can be expressed as a scaling factor of the ν_τ CC cross section as described in section 3.1; therefore, the measurement of the ν_τ normalisation can be interpreted as a direct constraint on the ν_τ CC cross section.

The hypothesis of the unitarity of the neutrino mixing matrix made in atmospheric neutrino experiments constrains the sum of the oscillation probabilities as follows

$$\sum_{\gamma} P_{\beta\gamma}(E, L) = 1 \quad (3.1)$$

where $P_{\beta\gamma}$ is the oscillation probability from the neutrino flavour β to γ , L the travelled distance, and E the energy. In this work, the effect of non-unitarity on atmospheric neutrino oscillations, which makes the sum of the measured oscillation probabilities deviate from 1, is studied. The efforts are concentrated on a single parameter that scales the τ -row of the PMNS matrix. This individual parameter affects neutrino oscillation probabilities in the ν_τ appearance but also in the ν_μ disappearance channel [22, 40], as described in section 3.2, enhancing the sensitivity of KM3NeT/ORCA. Section 3.2 describes the theoretical framework for this study involving the non-unitarity parametrisation of the neutrino mixing.

3.1 From the ν_τ normalisation to the ν_τ CC cross section measurement

The rate of detected charged-current $\nu_\tau/\bar{\nu}_\tau$ events can be expressed as

$$N_\tau = \int_{E_\nu} \int_{\cos(\theta)} \sigma_\tau^{\text{meas}} \{ \Phi_e P_{e\tau} + \Phi_\mu P_{\mu\tau} \} k_N \times \epsilon \times dE_\nu \times d\cos(\theta), \quad (3.2)$$

where $\sigma_\tau^{\text{meas}}$ is the measured ν_τ cross section, E_ν the neutrino energy, θ the zenith angle, k_N the number of target nucleons in the detector volume, ϵ the detector efficiency, and Φ_α the atmospheric neutrino flux, for any flavour α , computed with the Honda model [33]. The measured ν_τ cross section $\sigma_\tau^{\text{meas}}$ is related to the theoretical expectation value σ_τ^{th} used in the Monte Carlo simulation [31] through an energy independent scale factor S_τ , defined as

$$\sigma_\tau^{\text{meas}}(E_\nu) = S_\tau \times \sigma_\tau^{\text{th}}(E_\nu). \quad (3.3)$$

S_τ acts as a normalisation, which follows the naming convention used in [23].

3.2 Non-unitary neutrino mixing

Beyond Standard Model physics is required to explain the origin of neutrino masses. These may arise from a dimension-5 operator [41] that is generated in the seesaw model. It relies on the existence of Heavy Neutral Leptons (HNLs) as seesaw messengers to explain the lightness of neutrino masses compared to other fermions. In the high-scale scenario, where the HNL masses involved are above the electroweak symmetry breaking and beyond experimental reach, the heavy states are kinematically forbidden. In the low-scale scenario, like the inverse or linear seesaw [42], the neutrino mixing matrix can be written as a $n \times n$ ($n > 3$) unitary matrix. In this scenario, considered here, the new states are kinematically accessible and participate in the neutrino oscillation. The extended mixing matrix remains unitary but $\sum_{\gamma} P_{\beta\gamma}(E, L) \neq 1$ due to the contribution of oscillations involving HNLs.

The non-unitary neutrino mixing (NUNM) can be used to parametrise the effect of $n - 3$ HNLs on the neutrino oscillation. This includes the possibility of 3 heavy neutrinos as partners of the standard model (SM) neutrinos, that naturally arise in the type-1 seesaw model. The extended $n \times n$ unitary neutrino mixing matrix U can be decomposed as in [43], in a product of $n(n - 1)/2$ rotations:

$$U = R_{n-1\,n} \cdot R_{n-2\,n} \cdots R_{3\,n} \cdot R_{2\,n} \cdot R_{1\,n} \cdots R_{2\,3} \cdot R_{1\,3} \cdot R_{1\,2}, \quad (3.4)$$

with R_{ij} representing rotations of angle θ_{ij} with i and j between 1 and n , and $(R_{2\,3} \cdot R_{1\,3} \cdot R_{1\,2})^{3 \times 3} = U_{\text{PMNS}}$ the usual neutrino mixing matrix. When rotations are applied in this specific order, the non-unitary part of U is lower triangular. Based on that, the general structure of the non-unitarity parametrisation is constructed. The lower triangular matrix α is defined as

$$\alpha = \begin{pmatrix} \alpha_{11} & 0 & 0 \\ \alpha_{21} & \alpha_{22} & 0 \\ \alpha_{31} & \alpha_{32} & \alpha_{33} \end{pmatrix}, \quad (3.5)$$

involving 9 new parameters in addition to the SM mixing angles and phases with real diagonal parameters and complex non-diagonal ones. This lower triangular formalism was first introduced by Okubo in 1962 [44], in [20, 41, 45], and tested on data in [46]. The non-unitary matrix α applied to the unitary mixing matrix U_{PMNS} gives the non-unitary matrix N as

$$N = \alpha U_{\text{PMNS}}. \quad (3.6)$$

This model avoids the computationally intensive parametrisation of these effects with a $n \times n$ unitary mixing matrix in a sterile neutrino model where the number of corresponding angles and phases grows as n^2 . The Hamiltonian in matter $H_m^{n \times n}$ that describes the evolution of the neutrino state is defined as

$$H_m^{n \times n} = \Delta^{n \times n} + U^\dagger V^{n \times n} U \quad (3.7)$$

with $V^{n \times n} = \text{diag}\{V, 0, \dots, 0\}$ and $\Delta^{n \times n} = \text{diag}\{\Delta, \Delta m_{41}^2, \dots, \Delta m_{n1}^2\}$. Here, the effective matter potential V that describes the effect of the medium on the neutrino oscillation, and

the mass splitting matrix Δ are defined as

$$\Delta = \frac{1}{2E} \begin{bmatrix} 0 & 0 & 0 \\ 0 & \Delta m_{21}^2 & 0 \\ 0 & 0 & \Delta m_{31}^2 \end{bmatrix} \quad V = \begin{bmatrix} V_{CC} + V_{NC} & 0 & 0 \\ 0 & V_{NC} & 0 \\ 0 & 0 & V_{NC} \end{bmatrix}, \quad (3.8)$$

with the effective potential due to charged-current and neutral-current interactions $V_{CC} = \sqrt{2}G_F n_e$ and $V_{NC} = -\frac{1}{\sqrt{2}}G_F n_n$, respectively, G_F the Fermi constant, n_e and n_n the electron and neutron number density, respectively.

In this work, the mass of the heavy neutrino $m_{i>3} \gg \sqrt{\Delta m_{31}^2}$, therefore the model is not sensitive to the HNLs fast oscillations that are averaged out. As discussed in [46], under this approximation and in the mass basis, $H_m^{n \times n}$ can be truncated into a 3×3 matrix, therefore

$$H_m^{3 \times 3} = \Delta + N^\dagger V N. \quad (3.9)$$

Only the matter potential proportional to the electron density V_{CC} is relevant in the unitary case; indeed, the matter potential V_{NC} , proportional to the neutron density, affects uniformly all flavours, and is thus equivalent to an absolute phase shift of the neutrino state. Instead, the effect of V_{NC} on active neutrinos only, must be accounted for in the non-unitary case.

The calculation of the probability after propagation is affected by the non-unitary mixing matrix N . $P_{\beta\gamma}^\alpha$ describes the oscillation probability to observe flavour γ from initial flavour β after propagation of a distance L in a fixed density in the presence of unitarity violation (super script α) as

$$P_{\beta\gamma}^\alpha = \left| \left(N e^{-iH_m^{3 \times 3} L} N^\dagger \right)_{\gamma\beta} \right|^2. \quad (3.10)$$

The propagation through multiple density layers with non-unitary mixing is described in [47] and has been implemented in the OscProb software [35] used in this analysis.

The non-observation of very-short-distance oscillations [48, 49] and searches for light sterile neutrino in atmospheric neutrino experiments, yield the constraints (at 95% confidence level) reported in table 1 on the parameters of α , with α_{33} being the most weakly constrained parameter [50].

This parameter affects the ν_τ appearance, as it multiplies directly the third row of U_{PMNS} as well as other channels including the ν_μ disappearance, due to the V_{NC} term. The $V_{NC}=0$ case is included in this work, so that the results can be compared to those from [39], which uses the equivalent CC+NC ν_τ normalisation naming convention. In that specific case, the modified probability for a neutrino of initial flavor β to remain active is defined as

$$P_\beta^\alpha = P_{\beta e} + P_{\beta \mu} + \alpha_{33}^2 P_{\beta \tau}. \quad (3.11)$$

The modified ν_τ CC event rates is given by N_τ^α and NC event rates by N_{NC}^α . They are defined as $N_\tau^\alpha = \alpha_{33}^2 N_\tau$ and by

$$N_{NC}^\alpha = \int_E \int_{\cos(\theta)} \sigma_{NC} \left\{ \Phi_e[P_e^\alpha] + \Phi_\mu[P_\mu^\alpha] \right\} k_N \times \epsilon \times dE \times d\cos\theta. \quad (3.12)$$

Present bounds		
α_{33}	$>$	0.90
α_{22}	$>$	0.978
α_{11}	$>$	0.976
$ \alpha_{32} $	$<$	0.012
$ \alpha_{31} $	$<$	0.069
$ \alpha_{21} $	$<$	0.025

Table 1. Current bounds at the 95% CL derived from [50] for the NUNM parameters.

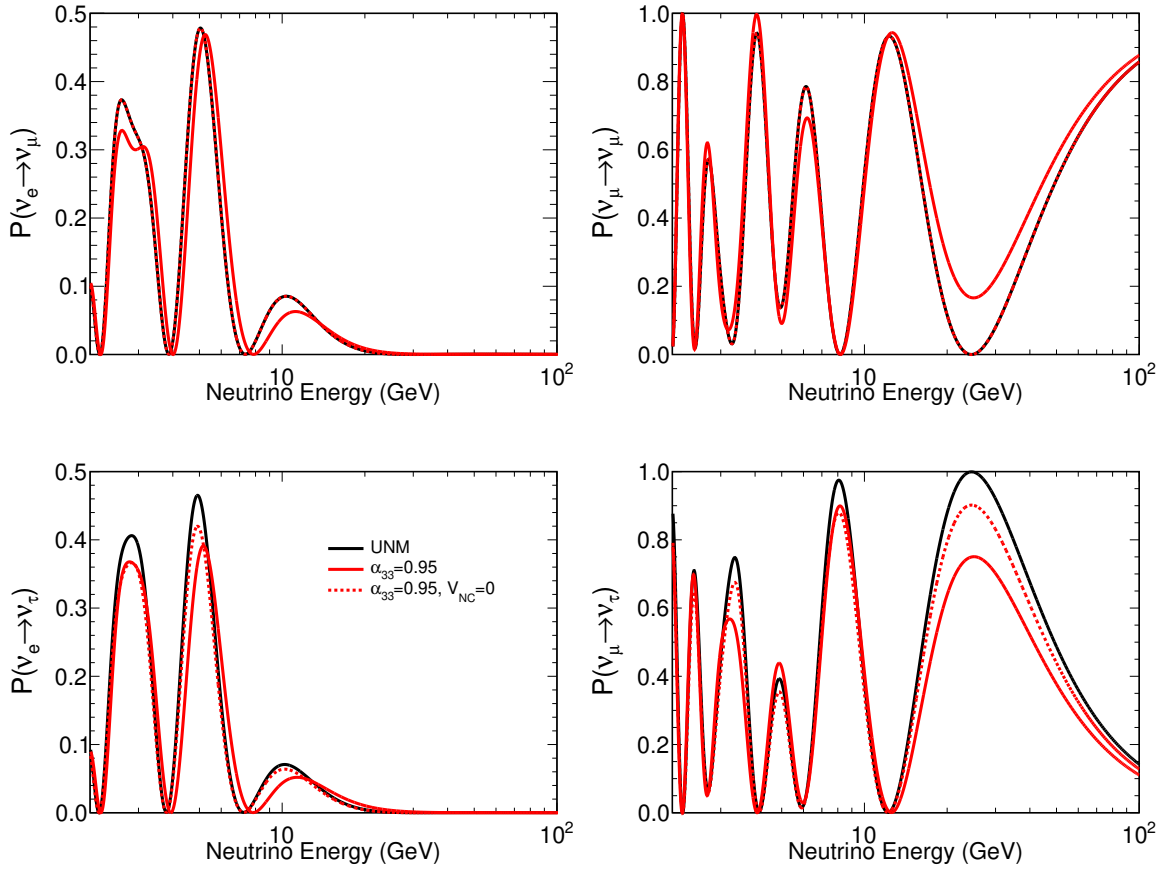


Figure 1. The oscillation probabilities are shown for vertically up-going ($\cos(\theta) = -1$) neutrinos as a function of neutrino energy for a non-unitary parameter α_{33} being non-zero and for $\sin(2\theta_{23}) = 1$ with and without including the neutral-current potential. The black curve illustrates the Unitary Neutrino Mixing (UNM) case in matter.

However, in matter, the equivalence between the non-unitarity framework and the CC+NC approach is broken and the difference in the effect is shown in figure 1. This illustrates that there is a significant contribution of the muon disappearance channel on constraining α_{33} , in addition to the tau appearance channel.

By measuring Earth-crossing neutrino oscillations, KM3NeT/ORCA can achieve a competitive sensitivity to α_{33} . In addition, when assuming $\cos^2 \theta_{34} = \alpha_{33}^2$, this non-unitarity test in the low-scale scenario can be translated into a 1-sterile neutrino case study as described in [51], under the hypothesis of $\Delta m_{41}^2 \gg \frac{E}{L}$, $\theta_{14} = \theta_{24} = 0$.

4 Event selection

The event selection aims at rejecting atmospheric muons and events due to environmental optical background as much as possible and classifying the remaining neutrino candidates according to their topology in a multi-step procedure that was optimised for the measurement of neutrino oscillation parameters. The procedure is summarised hereafter. More details can be found in [29] and references therein.

Events induced by optical background are rejected at first by requiring a minimum number of triggered hits and a good track-reconstruction likelihood. After this first selection, the event sample is mostly composed of atmospheric muons, that can be reduced by selecting only up-going reconstructed tracks, since atmospheric muons from below the horizon are absorbed by the Earth. Finally, a Boosted Decision Tree classifier is used to further discriminate atmospheric muons from atmospheric neutrinos. The remaining events are categorised according to their reconstructed event topology into tracks and showers. The track-like class is split into Low Purity Tracks and High Purity Tracks where purity refers to the contamination of atmospheric muons, that is 4% and 0.4%, respectively.

The final dataset contains 5828 events that are distributed across the different classes as shown in table 2. In addition, the expected number of events per interaction channel is presented. It is given by the MC event numbers from a fit (see section 5) under the assumption that $S_\tau = 1$ and $\alpha_{33} = 1$. According to this model, a total of 185 CC tau neutrino and antineutrino interactions are expected. The majority (77%) are selected in the shower class because τ leptons produced in ν_τ CC interactions decay into electrons or hadrons with a total branching ratio of about 83%. The remaining fraction of ν_τ CC interactions is evenly distributed between the High Purity Tracks and the Low Purity Tracks.

5 Analysis method

The ν_τ normalisation S_τ and the non-unitarity parameter α_{33} are measured by fitting a model to the observed event distributions. The event distributions predicted by the model depend on the parameter of interest (PoI), as well as on several nuisance parameters ϵ that account for systematic uncertainties.

Since neutrino oscillation probabilities depend on the path length L of the neutrino from its production to its detection and on the energy of the neutrino E , these quantities are estimated using reconstruction algorithms. The analysis is carried out in the two-dimensional space of reconstructed energy and cosine of the reconstructed zenith angle, which is proportional to

Event Type	Showers	High Purity Tracks	Low Purity Tracks	Total
ν_e CC + $\bar{\nu}_e$ CC	603	51	85	739
ν_μ CC + $\bar{\nu}_\mu$ CC	902	1777	1786	4465
ν_τ CC + $\bar{\nu}_\tau$ CC	143	22	20	185
ν NC + $\bar{\nu}$ NC	289	13	22	324
atmospheric μ + $\bar{\mu}$	22	7	89	118
Total MC fitted	1959	1870	2002	5831
Total MC before fit	2089	1845	1878	5812
Total Data	1958	1868	2002	5828

Table 2. Number of events and composition for the three classes for the MC fitted to the observed data, under the assumption that ν_τ normalisation $S_\tau = 1$ and non-unitarity parameter $\alpha_{33} = 1$ holds. For completeness, the expected event numbers calculated for oscillation parameters from [4] and all systematics at their nominal values are reported in the second to last row. The observed number of events is reported in the last row.

the neutrino path length. The two-dimensional distribution is built separately for each of the classes introduced in section 4. Ten equally-spaced bins are used for $\cos \theta$, in the range $[-1, 0]$ (where $\cos \theta = -1$ corresponds to vertically up-going neutrinos). For the reconstructed energy, 15 bins between 2 GeV and 1 TeV are used. The binning is defined in logarithmic scale and optimised such that the expected number of events in each bin is at least 2.

The different models introduced in section 3 are fitted to the data through the minimisation of a negative log-likelihood function $-2 \log \mathcal{L}$ defined as:

$$-2 \log \mathcal{L} = \sum_i 2 \left[\left(\beta_i N_i^{\text{mod}} - N_i^{\text{dat}} \right) + N_i^{\text{dat}} \ln \left(\frac{N_i^{\text{dat}}}{\beta_i N_i^{\text{mod}}} \right) \right] + \frac{(\beta_i - 1)^2}{\sigma_{\beta_i}^2} + \sum_k \left(\frac{\epsilon_k - \langle \epsilon_k \rangle}{\sigma_{\epsilon_k}} \right)^2 \quad (5.1)$$

The sum in the first term of equation (5.1) runs over each bin i of each of the aforementioned two-dimensional histograms, computing the Poisson likelihood of observing a number of events in data, N^{dat} , given a model with an expectation of N^{mod} . The sum in the second term of equation (5.1) runs over the nuisance parameters ϵ for which constraints have been set by other experiments. This knowledge is taken into account by fitting the systematic uncertainties ϵ with a Gaussian prior σ_ϵ in order to incorporate it in a consistent manner. The limited Monte Carlo statistics is addressed by introducing normally distributed coefficients β_i following the Barlow and Beeston light method [52, 53]. The minimisation is based on the MINUIT library [54]. A detailed description of the nuisance parameters is given in [29], and summarised in table 3. These also include oscillation parameters. The analysis is, however, only sensitive to Δm_{31}^2 and θ_{23} which are left unconstrained in the fit, while the other oscillation parameters are fixed to the NuFit v5.0 reference values [4]. Uncertainties on the shape of the atmospheric neutrino flux are represented by the uncertainty on the spectral index δ_γ and on the slope between horizontal and vertically up-going neutrinos δ_θ . In addition, the ratios between neutrinos and antineutrinos $s_{\mu\bar{\mu}}$, $s_{e\bar{e}}$ of muon and electron neutrinos respectively, as well as the flavour ratio $s_{\mu e}$ are the uncertainties

Parameter	Constraints
θ_{23}	unconstrained
Δm_{31}^2	unconstrained
f_{all}	unconstrained
f_{HPT}	unconstrained
f_{S}	unconstrained
f_{HE}	50%
f_{μ}	unconstrained
f_{NC}	20%
$s_{\mu\bar{\mu}}$	5%
$s_{e\bar{e}}$	7%
$s_{\mu e}$	2%
δ_{γ}	0.3
δ_{θ}	2%
E_s	9%
S_{τ}	depending on analysis

Table 3. All nuisance parameters and their treatment in the fit. A description of the individual parameters can be found in the text.

in the flavour composition of the atmospheric neutrino flux. All systematic uncertainties related to the atmospheric neutrino flux are fitted with constraints considering motivations described in [55]. A nuisance parameter f_{NC} is introduced to account for uncertainties in the all-flavour neutral-current interaction rate. The PoI for the cross section measurement, S_{τ} , is an additional nuisance parameter in the non-unitarity fit. The parameter f_{HE} addresses the differences between the light simulation approaches that have been used in the simulation of low- and high-energy neutrinos, while the energy scale E_s combines several uncertainties associated with the detector response such as PMT efficiencies and water properties extracted from [26]. Finally, four normalisation nuisance parameters are considered: overall normalisation f_{all} , high-purity track normalisation f_{HPT} , shower normalisation f_{S} , and atmospheric muon normalisation f_{μ} , allowing for a scaling of the corresponding event numbers. Regarding the cross section, uncertainties related to quasi-elastic scattering and resonance production are found to have a negligible impact on this study while uncertainties related to the deep inelastic scattering channel are fully accommodated by the implemented systematics such as channel-dependent normalisations and δ_{γ} .

The sensitivity is estimated by profiling the log-likelihood ratio (LLR)

$$\text{LLR} = -2\Delta \log \mathcal{L} = \min_{\epsilon} (-2 \log \mathcal{L}) - \min_{\text{PoI}, \epsilon} (-2 \log \mathcal{L}) \quad (5.2)$$

where the first term represents the likelihood evaluated using the best-fit values of the nuisance parameters and a fixed value for the PoI, and the second term represents the

Model	PoI	N_{ν_τ} CC	N_ν NC
S_τ	$0.48^{+0.5}_{-0.33}$	92^{+90}_{-63}	340^{+16}_{-16}
α_{33}	$0.993^{+0.026}_{-0.025}$	170^{+5}_{-9}	325^{+1}_{-4}
α_{33} ($V_{\text{NC}} = 0, S_\tau = 1$)	$0.83^{+0.20}_{-0.25}$	132^{+60}_{-61}	313^{+11}_{-13}

Table 4. Measured ν_τ normalisation S_τ and non-unitarity parameter α_{33} and corresponding number N_{ν_τ} CC of charged-current tau neutrino and N_ν NC of neutral-current all flavour neutrino interactions for the different considered models. The 1σ uncertainties for S_τ and α_{33} are obtained with the FC approach described in section 5, whereas in case of α_{33} ($V_{\text{NC}} = 0, S_\tau = 1$), they are based on Wilks’ theorem.

likelihood evaluated using the best-fit values for the nuisance parameters and for the PoI. The LLR is therefore evaluated for different values of S_τ or α_{33} to compute the sensitivity and confidence intervals.

The confidence intervals are evaluated with the Feldman-Cousins (FC) method [56] generating 1000 pseudo experiments for several values of the PoI. The nuisance parameters are set to their nominal values assuming normal ordering. Poissonian fluctuations are added to the bin contents in order to account for statistical fluctuations. Finally, the constraints on the nuisance parameters for the fit are randomised by drawing them from a Gaussian distribution according to the mean and uncertainty given in table 3 [57]. The uncertainty on the limits is obtained by exploiting a bootstrapping technique [58].

6 Results

The analysis procedures described in section 5 have been applied to fit each of the models described in section 3 to the data set described in section 4. The results are summarised in tables 4 and 5, which respectively show the best-fit values for the parameters of interest of each model, and the best-fit values for the nuisance parameters. The PoI is S_τ when measuring the ν_τ CC cross section and α_{33} in case of probing non-unitarity.

6.1 ν_τ CC cross section measurement

The parameter S_τ and the corresponding 1σ uncertainty were measured to be $0.48^{+0.5}_{-0.33}$. The best fit corresponds to a total number of 92^{+90}_{-63} observed ν_τ CC events and a total of 340^{+16}_{-16} NC interactions. The profiled log-likelihood for S_τ reported in figure 2 allows for the visualisation of the best-fit value as well as the FC correction of the 68 % and 90 % confidence intervals. The dashed lines indicate the 68% and 90% CLs according to Wilks’ theorem [59], whereas the coloured lines represent the CLs from the FC correction with their corresponding 1σ uncertainty derived from bootstrapping. As expected, the CLs from Wilks’ theorem and FC mostly deviate from each other at low values of S_τ , i.e. close to the physical boundary. Above $S_\tau = 0.25$ (0.50) the difference for the 68% (90%) CL is negligible. The p-value to exclude the hypothesis of $S_\tau = 0$ is $(5.9 \pm 0.8)\%$, calculated as the fraction of pseudo datasets for which $-2\Delta \log \mathcal{L}$ is larger than the observed one.

The L/E distribution is shown in figure 3. Even though it is not used in the analysis, this distribution helps to visualise the results since the oscillation probability is a function

	S_τ	α_{33}
Systematic uncertainty	$\epsilon \pm 1\sigma$	$\epsilon \pm 1\sigma$
$\theta_{23} [^\circ]$	46^{+4}_{-4}	46^{+4}_{-4}
$\Delta m_{31}^2 [\times 10^{-3} \text{ eV}^2]$	$2.15^{+0.29}_{-0.28}$	$2.18^{+0.19}_{-0.35}$
f_{all}	$1.09^{+0.17}_{-0.11}$	$1.11^{+0.11}_{-0.12}$
f_{HPT}	$0.91^{+0.04}_{-0.04}$	$0.92^{+0.04}_{-0.04}$
f_{S}	$0.92^{+0.06}_{-0.06}$	$0.92^{+0.06}_{-0.06}$
f_{HE}	$1.50^{+0.33}_{-0.30}$	$1.59^{+0.32}_{-0.29}$
f_μ	$0.5^{+0.44}_{-0.4}$	$0.51^{+0.4}_{-0.35}$
f_{NC}	$0.89^{+0.20}_{-0.20}$	$0.86^{+0.19}_{-0.19}$
$s_{\mu\bar{\mu}}$	$0.00^{+0.05}_{-0.05}$	$0.00^{+0.05}_{-0.05}$
$s_{e\bar{e}}$	$0.01^{+0.07}_{-0.07}$	$0.01^{+0.07}_{-0.07}$
$s_{\mu e}$	$-0.004^{+0.020}_{-0.020}$	$-0.004^{+0.020}_{-0.019}$
δ_γ	$-0.019^{+0.027}_{-0.026}$	$-0.019^{+0.025}_{-0.025}$
δ_θ	$-0.002^{+0.019}_{-0.019}$	$-0.005^{+0.019}_{-0.019}$
E_s	$0.98^{+0.11}_{-0.08}$	$1.03^{+0.08}_{-0.11}$
S_τ	PoI	$0.92^{+0.19}_{-0.18}$

Table 5. All systematic uncertainties and their best-fit values along with their 1σ post-fit uncertainties. The description of the parameters is given in the text.

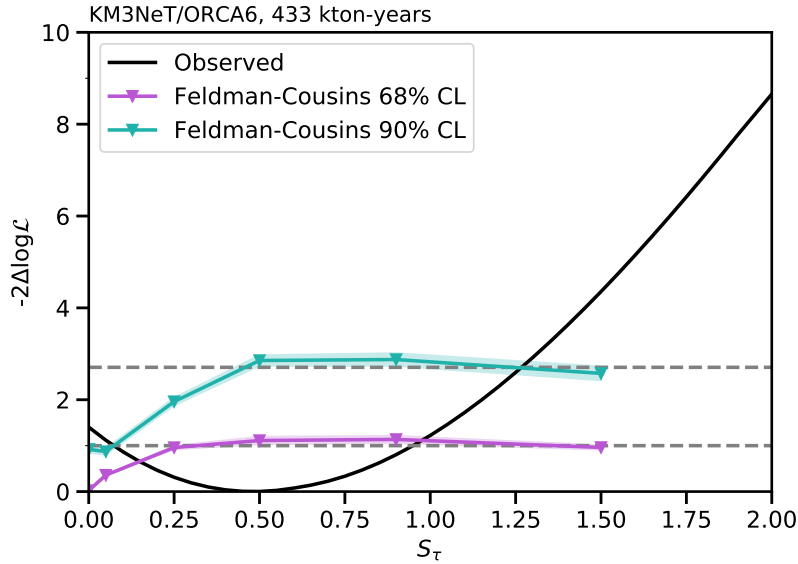


Figure 2. Measured log-likelihood profile with Feldman-Cousins correction for the model probing the ν_τ CC cross section (black line). The horizontal dashed lines represent the 68% and 90% CL thresholds according to Wilks' theorem, whereas the coloured lines indicate the Feldman-Cousins correction with the corresponding 1σ uncertainty.

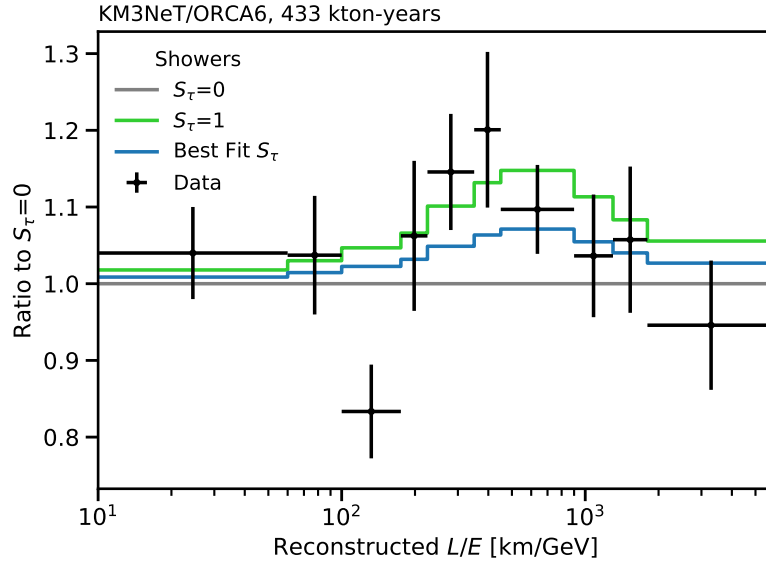


Figure 3. Measured L/E distribution for the shower class (black crosses) with best-fit model (blue line) and a model with nominal ν_τ normalisation S_τ (green line) with respect to no- ν_τ appearance.

of L/E . The ratio for the observed data compared to the best-fit model and a model with $S_\tau = 1$ to a model with the extreme case $S_\tau = 0$ is presented. The distribution is shown for the shower class, that is the most sensitive class to the observation of ν_τ appearance.

The measured ν_τ CC cross section is related to the theoretical expectation following equation (3.3). In figure 4, the energy-dependent theoretical expectation from [31] is scaled using the best-fit value $S_\tau = (0.48^{+0.5}_{-0.33})$, while the median of the ν_τ CC true energy distribution E_{ν_τ} is at 20.3 GeV between 12.3 and 35.9 GeV at 68% CL. The corresponding theoretical expectation is $\sigma_\tau^{\text{th}} = 5.29 \times 10^{-38} \text{ cm}^2$, calculated as the average of the neutrino and antineutrino cross section according to equation (6.1)

$$\begin{aligned}
 \sigma_\tau^{\text{th}} &= r_\nu \sigma_\tau + (1 - r_\nu) \sigma_{\bar{\tau}} \\
 r_\nu &= X_\nu / (X_\nu + \lambda X_{\bar{\nu}}) \\
 X_\nu &= \int_{E_\nu} \int_{\cos(\theta)} \sigma_\tau \{ \Phi_e P_{e\tau} + \Phi_\mu P_{\mu\tau} \} \times dE_\nu \times d\cos(\theta) \\
 X_{\bar{\nu}} &= \int_{E_{\bar{\nu}}} \int_{\cos(\theta)} \sigma_{\bar{\tau}} \{ \Phi_{\bar{e}} P_{\bar{e}\bar{\tau}} + \Phi_{\bar{\mu}} P_{\bar{\mu}\bar{\tau}} \} \times dE_{\bar{\nu}} \times d\cos(\theta).
 \end{aligned} \tag{6.1}$$

The neutrino fraction is found as $r_\nu = 0.72$. X_ν and $X_{\bar{\nu}}$ represent — up to some constants which cancel in the ratio — the neutrino and antineutrino interaction rate while λ accounts for potential differences in the reconstruction efficiencies between neutrino and antineutrino events due to their non-identical interaction inelasticities. However, the invisible neutrino from the tau decay fully compensates for the presumed inelasticity effect and $\lambda = 1.0$ is derived. Hence, the measured cross section is $\sigma_\tau^{\text{meas}} = (2.5^{+2.6}_{-1.8}) \times 10^{-38} \text{ cm}^2$.

The effect of the nuisance parameters on the fit result of S_τ is presented in figure 5, where blue bars represent the impact of a nuisance parameter on S_τ . Each of them is calculated by

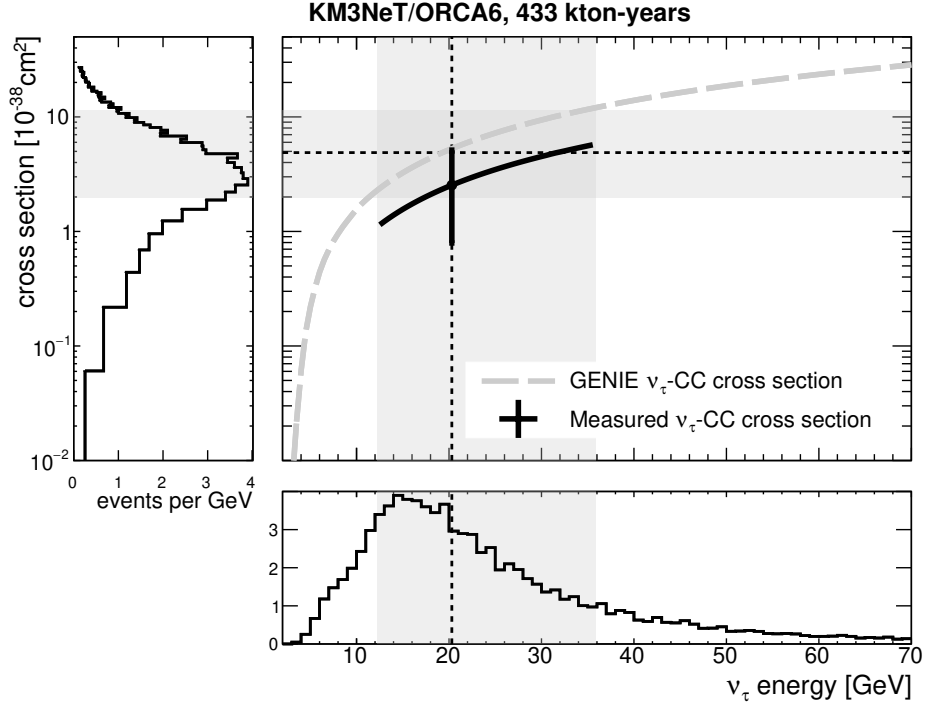


Figure 4. Measurement of the charged-current tau neutrino cross section in black, compared to the $\nu_\tau/\bar{\nu}_\tau$ averaged theoretical expectation in grey as a function of the energy. The light grey bands represent the 68% range of the distributions.

performing two fits: one in which the nuisance parameter is fixed to the best-fit value from table 5 plus its 1σ post-fit uncertainty and one in which it is fixed to the best-fit value minus its uncertainty. Each fit yields a new fitted value for S_τ denoted as S_τ^{shift} , which allows for the calculation of the impact as $(S_\tau^{\text{shift}} - S_\tau^{\text{bf}})/\sigma_{S_\tau}$, where S_τ^{bf} is the best-fit value of S_τ , and σ_{S_τ} its 1σ error. This plot also shows the pulls for each nuisance parameter (black markers), which are defined as the difference between the best-fit value ϵ_{BF} and the expected value ϵ_{CV} of the corresponding parameter, divided by its uncertainty $(\epsilon_{\text{BF}} - \epsilon_{\text{CV}})/\sigma$. If available, σ is given by the pre-fit uncertainty (see table 3), and the error bars represent the ratio of the post-fit and pre-fit uncertainties $\sigma_\epsilon^{\text{post-fit}}/\sigma_\epsilon^{\text{pre-fit}}$. For the parameters that are fitted without constraints, σ is given by the post-fit uncertainty, and the error bars are set to 1.

The high-energy light simulation f_{HE} and the f_{NC} normalisation are the nuisance parameters with the largest impact on S_τ . The energy scale E_s , i.e. the detector-related uncertainties, has a non-negligible impact. Regarding the atmospheric neutrino flux parameters, only the slope between up- and horizontally-going neutrinos δ_θ has a sizeable influence, whereas the three composition ratios do not affect the S_τ measurement significantly. Systematic uncertainties that are fitted without constraints have larger pulls, as expected. Among the systematic uncertainties that are fitted with a constraint, only f_{HE} and f_{NC} show large pulls. Except for δ_γ , for which the post-fit uncertainty is significantly reduced compared to the pre-fit one, all other nuisance parameters have a negligible improvement in their post-fit uncertainty.

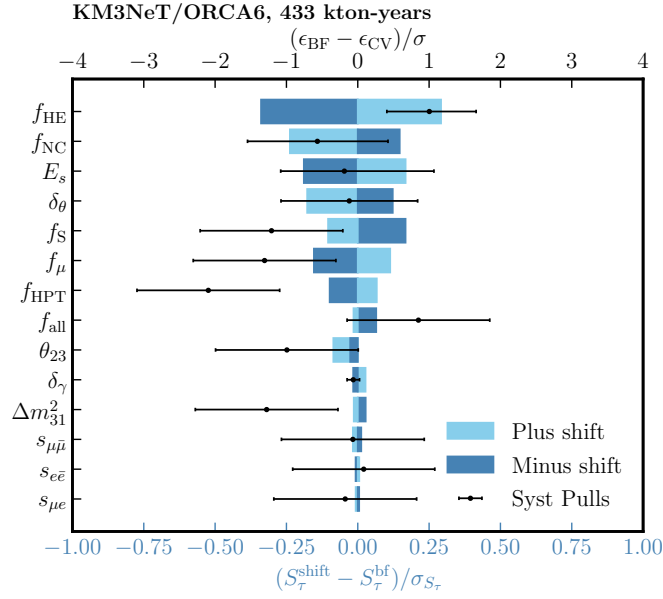


Figure 5. Impact on the best-fit value of S_τ when shifting a systematic uncertainty by plus (light blue) or minus (dark blue) the 1σ post-fit uncertainty from its best-fit value (lower x -axis). The pulls (black markers, upper x -axis) represent the deviation of the best-fit value from the central value of the parameter with respect to its uncertainty.

6.2 Probing unitarity

The scale factor for the ν_τ CC cross section S_τ is included as a nuisance parameter for the measurement of α_{33} , and assigned a 20 % uncertainty. The best-fit value is $\alpha_{33} = 0.993^{+0.026}_{-0.025}$ corresponding to 170^{+5}_{-9} ν_τ CC and a total of 325^{+1}_{-4} NC events. The profiled log-likelihood ratio is presented in figure 6. On the left hand side, it is compared to a model without neutral-current potential, i.e. $V_{\text{NC}} = 0$, and with $S_\tau = 1$ for which the best-fit value is $\alpha_{33} = 0.83^{+0.20}_{-0.25}$, corresponding to 132^{+60}_{-61} ν_τ CC and a total of 313^{+11}_{-13} NC events. The latter results can be compared to the results published by the IceCube collaboration [39] following the physics interpretation from section 3, where the CC+NC ν_τ normalisation corresponds to α_{33}^2 with $V_{\text{NC}} = 0$. The difference in the sensitivities for both models originates from the fact that when $V_{\text{NC}} = 0$, α_{33} affects only the ν_τ appearance probability, as shown in figure 1. If on the other hand $V_{\text{NC}} \neq 0$, α_{33} additionally affects other oscillation channels including ν_μ disappearance leading to better constraints (as discussed in section 7).

The FC correction is shown on the right side of figure 6. Due to the high computational cost, it is only applied for the case where S_τ is fitted and V_{NC} is included. The results are similar to Wilks' theorem except for values of α_{33} that are close to 1, which suggests that there is a physical boundary. In fact, $\alpha_{33} > 1$ is an unphysical region for which the sum of oscillation probabilities would exceed 1. Overall, the 68% confidence interval is almost unaffected, whereas the 95% confidence interval is slightly reduced with $\alpha_{33} > 0.95$ at 95% CL.

The best fit is consistent with unitarity, i.e. $\alpha_{33} = 1$ (p-value of $68\% \pm 2\%$), also when neglecting the V_{NC} potential and fixing the ν_τ CC cross section (0.6σ according to Wilks' theorem).

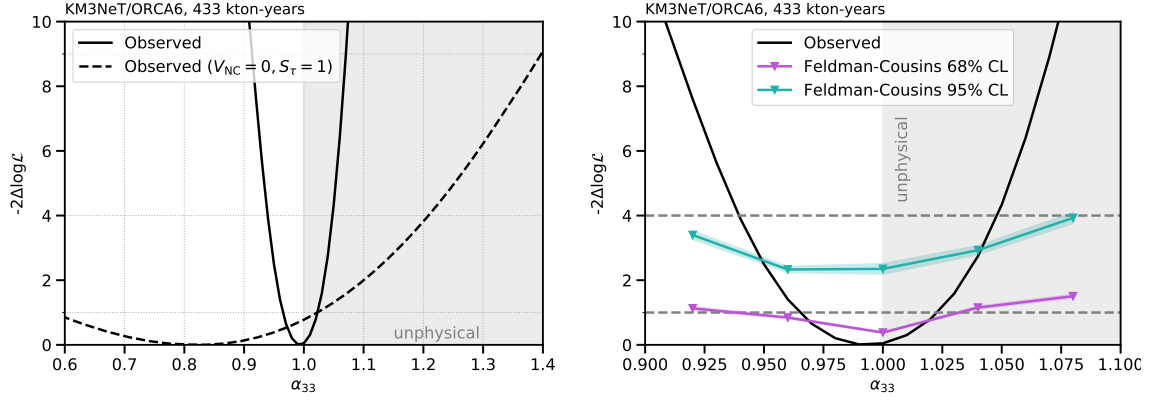


Figure 6. Measured log-likelihood profile for the model probing non-unitarity (black solid lines). *Left:* comparison with a model with $V_{\text{NC}} = 0$ and $S_\tau = 1$ (black dashed line). *Right:* Feldman-Cousins correction. The horizontal dashed lines represent the 68% and 95% CL thresholds according to Wilks’ theorem, whereas the coloured lines indicate the Feldman-Cousins correction with the corresponding 1σ uncertainty.

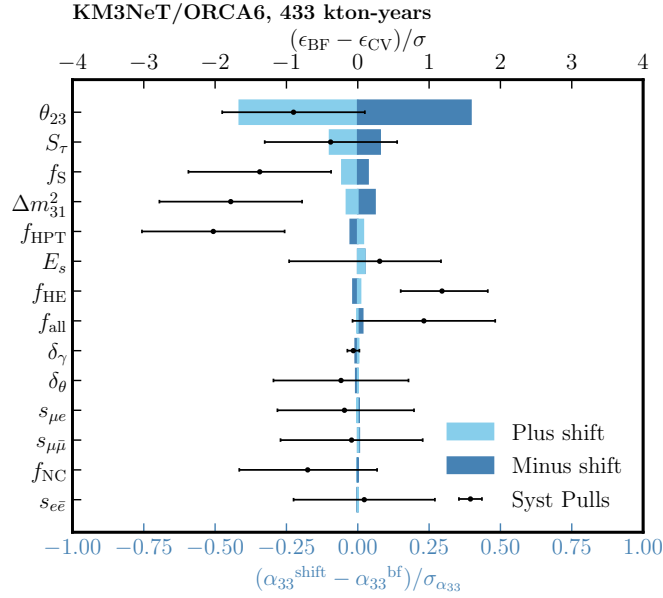


Figure 7. Impact on the best-fit value of α_{33} when shifting a systematic uncertainty by plus (light blue) or minus (dark blue) the 1σ post-fit uncertainty from its best-fit value (lower x -axis). The pulls (black markers, upper x -axis) represent the deviation of the best-fit value from the central value of the parameter with respect to its uncertainty.

Impacts and pulls for the measurement of α_{33} are presented in figure 7. Overall, the pulls of the nuisance parameters are comparable with what is obtained from the measurement of S_τ , except for E_s which has the opposite sign and a larger pull for Δm_{31}^2 . The nuisance parameter with the largest impact on α_{33} is clearly θ_{23} . This can be explained by the fact that both parameters affect the ν_μ disappearance. Additionally, α_{33} has an influence on the ν_τ appearance channel which can partly be compensated by S_τ leading to a non-negligible

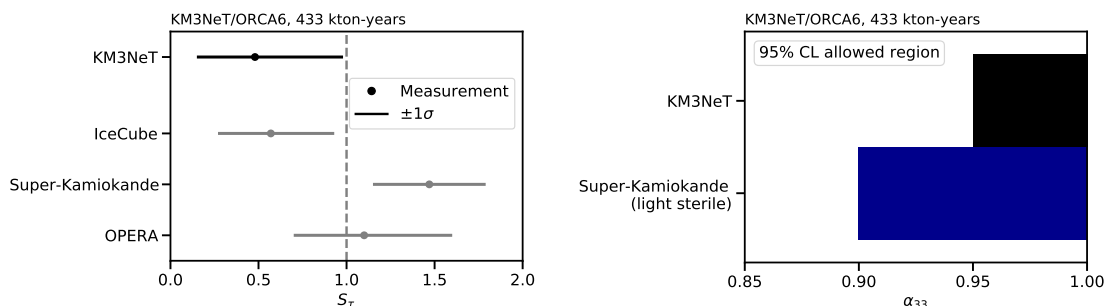


Figure 8. *Left:* comparison of S_τ measurement of ORCA6 with previous results on ν_τ CC cross section normalisation reported by OPERA [60], Super-Kamiokande [23], and IceCube [39]. *Right:* comparison of the α_{33} constraints of ORCA6 with the measurement from Super-Kamiokande derived in [50].

impact of this parameter. Apart from Δm_{31}^2 and f_S , the remaining nuisance parameters do not have a major impact on the result.

7 Comparison with other experiments and conclusions

Exploring ORCA6 data, two different analyses for the ν_τ normalisation and a test of the unitarity of the neutrino mixing matrix have been presented. A data sample of 433 kton-years exposure collected with 5% of the nominal instrumented volume (corresponding to 6 operational detection units) has been used for these analyses.

The first analysis explores the ν_τ appearance channel for a measurement of the ν_τ normalisation in the oscillation channel and explicitly assumes that the 3ν -paradigm is valid. Consequently, possible deviations of the ν_τ normalisation from 1 correspond to a scaling of the ν_τ CC event rate appearing from the oscillated atmospheric neutrino flux. This method is the same as the one used by other experiments exploring the ν_τ appearance oscillation either in the ν_μ CNGS beam (as in OPERA [19]) or in the atmospheric flux (as in Super-Kamiokande [23] and IceCube [39]). Nonetheless, the comparison among previously reported measurements is not straightforward because it is based on the detection of neutrinos oscillating at various neutrino energies and interacting in different media. Moreover, they have complementary sensitivity to the oscillation parameters; consequently, their treatment of the ν_τ normalisation fit is different. As in the atmospheric neutrino experiments Super-Kamiokande and IceCube, the ORCA6 analysis is performed on a statistical basis at the oscillation maximum around 25 GeV, due to the impossibility of directly tagging ν_τ and $\bar{\nu}_\tau$ events. In the ORCA6 event sample, 92_{-63}^{+90} ν_τ CC are observed. The ν_τ normalisation is fitted while keeping the θ_{23} angle and the Δm_{31}^2 mass splitting parameters free. Both normal and inverted ordering hypotheses as well as both θ_{23} octants are tested. In this way, the ν_τ normalisation is measured at $0.48_{-0.33}^{+0.5}$. The result is reported in terms of the significance of excluding the hypothesis $S_\tau = 0$, corresponding to a p-value of $(5.9 \pm 0.8) \%$. The up to date overview of the ν_τ CC normalisation measurements at 68% CL, including the ORCA6 result, is shown in the left plot of figure 8. None of the reported results shows statistical evidence of rejecting the standard 3ν -paradigm description. The achieved ORCA6 precision is already in line with other reported measurements, despite the reduced exposure of this analysis.

Experiment	Interaction medium	Energy [Gev]	N. of observed ν_τ	ν_τ CC cross section [nucleon $^{-1}$ 10^{-38} cm 2]
OPERA [60]	lead	≤ 20	10	$2.46^{+1.15}_{-0.98}$
Super-Kamiokande [23]	water	~ 25	338.1 ± 72.7	0.94 ± 0.20
ORCA6 (this work)	water	$20.3^{+8.0}_{-15.6}$	92^{+90}_{-63}	$2.5^{+2.6}_{-1.8}$

Table 6. Summary of the reported measurements of the ν_τ CC cross section and the corresponding statistics; the measurement from OPERA [60] has been converted to the same units as the results by Super-Kamiokande and ORCA6, dividing by the number of nucleons, depending on the target composition.

Under the hypothesis that the PMNS matrix is unitary, the ν_τ normalisation S_τ can be interpreted as a scale factor of the ν_τ CC cross section, allowing to constrain its measurement. Up to now, the uncertainty on the cross section is still one of the main sources of systematics in neutrino oscillation experiments, and measurements of the ν_τ CC cross section have been reported only by a few experiments. Hence, the diverse energy range and interaction media in the various experiments as well as the different impacts of systematic effects and uncertainties is an additional motivation for pursuing such measurements. In order to measure ν_τ CC cross section in ORCA6 data, the theoretical ν_τ cross section at the median of the ν_τ energy distribution, 20.3 GeV, within [12.3, 35.9] GeV at 68% CL, is considered. In the case of OPERA [60], the accessible neutrino energy range is restricted below 20 GeV and ν_τ CC interactions take place in a lead-based target. In experiments exploiting the ν_τ appearance channel from the atmospheric neutrino oscillation, the separation between ν_τ and $\bar{\nu}_\tau$ cannot be distinguished. Similarly to Super-Kamiokande, water is the interaction medium also for the ORCA6 measurement. Here, the ν_τ CC cross section is found to be $\sigma_{\text{meas}} = (2.5^{+2.6}_{-1.8}) \times 10^{-38}$ cm 2 and, as in the case of Super-Kamiokande, is still dominated by statistical uncertainties. This result is in good agreement with what was reported by OPERA. For a better comparison highlighting the difference among the experiments mentioned above, a summary of the ν_τ CC cross section measurements and the corresponding statistics is reported in table 6.

The second major result obtained from exploring this data sample is the explicit non-unitarity test of the neutrino mixing matrix, with the purpose of reaching a higher precision in validating the 3ν -paradigm. In the analysis described in this paper, the extension to a general $n \times n$ neutrino mixing matrix has been considered, by using the formalism described in section 3.2. Following this approach, the small non-zero neutrino masses are further suppressed by the existence of Heavy Neutral Leptons participating in the neutrino oscillations. Hence, the non-unitarity of the PMNS matrix is tested by extending it with additional parameters α_{ij} . A precise measurement of these extra parameters has been performed only partially on data in the accessible energy range of long-baseline neutrino experiments at accelerators [46]. The ORCA6 result is the first one obtained with Earth-crossing atmospheric neutrinos.

By exploring this neutrino sample, the best sensitivity on the most weakly constrained α_{33} parameter is accessible due to matter effects and the non-negligible contribution of the V_{NC} potential to oscillation probability. The fit of the α_{33} parameter is performed under two different assumptions: the most general one where S_τ is considered as a nuisance parameter with 20 % assigned uncertainty, and the specific case with $S_\tau=1$ and $V_{\text{NC}}=0$. The latter corresponds to a description where optimal knowledge on the ν_τ CC cross section is assumed and the neutron density does not contribute to the matter potential. This restricted description of the NUNM formalism is considered in the ORCA6 analysis for a direct comparison with the CC+NC normalisation reported by the IceCube Collaboration [39], under the PMNS unitarity hypothesis, where a variation in both CC and NC rates would affect the ν_τ normalisation measurement. In ORCA6, α_{33} is found at $0.83^{+0.20}_{-0.25}$; both the ν_τ CC and NC rates vary so that 132^{+60}_{-61} ν_τ CC and 313^{+11}_{-13} NC events are counted, respectively. This result is consistent with the IceCube CC+NC ν_τ normalisation measurement, and comparable precision is achieved in both analyses.

However, as described in section 3, the V_{NC} potential cannot be neglected in matter, given that it affects all the oscillation channels; hence, the α_{33} parameter is fitted including the V_{NC} potential and allowing the S_τ parameter to vary. Under these hypotheses, α_{33} is $0.993^{+0.026}_{-0.025}$. The reported ORCA6 limit on the α_{33} parameter, reduced at [0.95, 1.04], in the NUNM general formalism, is shown on the right plot of figure 8 and compared with the SuperKamiokande analysis. The limit on the α_{33} parameter is derived from [50]; each coloured band in the plot represents the 95% allowed region. The α_{33} limit found in this analysis is the most precise obtained to date. With increasing statistics being collected in larger detector geometries towards the nominal one, further improvements are expected.

Acknowledgments

The authors acknowledge the financial support of: KM3NeT-INFRADEV2 project, funded by the European Union Horizon Europe Research and Innovation Programme under grant agreement No 101079679; Funds for Scientific Research (FRS-FNRS), Francqui foundation, BAEF foundation. Czech Science Foundation (GAČR 24-12702S); Agence Nationale de la Recherche (contract ANR-15-CE31-0020), Centre National de la Recherche Scientifique (CNRS), Commission Européenne (FEDER fund and Marie Curie Program), LabEx UnivEarthS (ANR-10-LABX-0023 and ANR-18-IDEX-0001), Paris Île-de-France Region, Normandy Region (Alpha, Blue-waves and Neptune), France, The Provence-Alpes-Côte d’Azur Delegation for Research and Innovation (DRARI), the Provence-Alpes-Côte d’Azur region, the Bouches-du-Rhône Departmental Council, the Metropolis of Aix-Marseille Provence and the City of Marseille through the CPER 2021-2027 NEUMED project, The CNRS Institut National de Physique Nucléaire et de Physique des Particules (IN2P3); Shota Rustaveli National Science Foundation of Georgia (SRNSFG, FR-22-13708), Georgia; This work is part of the MuSES project which has received funding from the European Research Council (ERC) under the European Union’s Horizon 2020 Research and Innovation Programme (grant agreement No 101142396). This work was supported by the European Research Council, ERC Starting grant *MessMapp*, under contract no. 949555. The General Secretariat of Research and Innovation (GSRI), Greece; Istituto Nazionale di Fisica Nucleare (INFN) and Ministero

dell’Università e della Ricerca (MUR), through PRIN 2022 program (Grant PANTHEON 2022E2J4RK, Next Generation EU) and PON R&I program (Avviso n. 424 del 28 febbraio 2018, Progetto PACK-PIR01 00021), Italy; IDMAR project Po-Fesr Sicilian Region az. 1.5.1; A. De Benedittis, W. Idrissi Ibnsalih, M. Bendahman, A. Nayerhoda, G. Papalashvili, I. C. Rea, A. Simonelli have been supported by the Italian Ministero dell’Università e della Ricerca (MUR), Progetto CIR01 00021 (Avviso n. 2595 del 24 dicembre 2019); KM3NeT4RR MUR Project National Recovery and Resilience Plan (NRRP), Mission 4 Component 2 Investment 3.1, Funded by the European Union — NextGenerationEU, CUP I57G21000040001, Concession Decree MUR No. n. Prot. 123 del 21/06/2022; Ministry of Higher Education, Scientific Research and Innovation, Morocco, and the Arab Fund for Economic and Social Development, Kuwait; Nederlandse organisatie voor Wetenschappelijk Onderzoek (NWO), the Netherlands; The grant “AstroCeNT: Particle Astrophysics Science and Technology Centre”, carried out within the International Research Agendas programme of the Foundation for Polish Science financed by the European Union under the European Regional Development Fund; The program: “Excellence initiative-research university” for the AGH University in Krakow; The ARTIQ project: UMO-2021/01/2/ST6/00004 and ARTIQ/0004/2021; Ministry of Research, Innovation and Digitalisation, Romania; Slovak Research and Development Agency under Contract No. APVV-22-0413; Ministry of Education, Research, Development and Youth of the Slovak Republic; MCIN for PID2021-124591NB-C41, -C42, -C43 and PDC2023-145913-I00 funded by MCIN/AEI/10.13039/501100011033 and by “ERDF A way of making Europe”, for ASFAE/2022/014 and ASFAE/2022 /023 with funding from the EU NextGenerationEU (PRTR-C17.I01) and Generalitat Valenciana, for Grant AST22_6.2 with funding from Consejería de Universidad, Investigación e Innovación and Gobierno de España and European Union — NextGenerationEU, for CSIC-INFRA23013 and for CNS2023-144099, Generalitat Valenciana for CIDEAGENT/2018/034, /2019/043, /2020/049, /2021/23, for CIDEIG/2023/20, for CIPROM/2023/51 and for GRISOLIAP/2021/192 and EU for MSC/101025085, Spain; Khalifa University internal grants (ESIG-2023-008, RIG-2023-070 and RIG-2024-047), United Arab Emirates; The European Union’s Horizon 2020 Research and Innovation Programme (ChETEC-INFRA — Project no. 101008324).

Data Availability Statement. This article has no associated data or the data will not be deposited.

Code Availability Statement. This article has no associated code or the code will not be deposited.

Open Access. This article is distributed under the terms of the Creative Commons Attribution License ([CC-BY4.0](https://creativecommons.org/licenses/by/4.0/)), which permits any use, distribution and reproduction in any medium, provided the original author(s) and source are credited.

References

- [1] SUPER-KAMIOKANDE collaboration, *Evidence for oscillation of atmospheric neutrinos*, [*Phys. Rev. Lett.* **81** \(1998\) 1562](#) [[hep-ex/9807003](#)] [[INSPIRE](#)].
- [2] B. Pontecorvo, *Mesonium and antimesonium*, *Sov. Phys. JETP* **6** (1958) 429 [[INSPIRE](#)].

- [3] Z. Maki, M. Nakagawa and S. Sakata, *Remarks on the unified model of elementary particles*, *Prog. Theor. Phys.* **28** (1962) 870 [INSPIRE].
- [4] I. Esteban et al., *The fate of hints: updated global analysis of three-flavor neutrino oscillations*, *JHEP* **09** (2020) 178 [arXiv:2007.14792] [INSPIRE].
- [5] F. Capozzi et al., *Global constraints on absolute neutrino masses and their ordering*, *Phys. Rev. D* **95** (2017) 096014 [Addendum *ibid.* **101** (2020) 116013] [arXiv:2003.08511] [INSPIRE].
- [6] P.F. de Salas et al., *2020 global reassessment of the neutrino oscillation picture*, *JHEP* **02** (2021) 071 [arXiv:2006.11237] [INSPIRE].
- [7] NOvA collaboration, *Expanding neutrino oscillation parameter measurements in NOvA using a Bayesian approach*, *Phys. Rev. D* **110** (2024) 012005 [arXiv:2311.07835] [INSPIRE].
- [8] T2K collaboration, *Measurements of neutrino oscillation parameters from the T2K experiment using 3.6×10^{21} protons on target*, *Eur. Phys. J. C* **83** (2023) 782 [arXiv:2303.03222] [INSPIRE].
- [9] SNO collaboration, *Measurement of the rate of $\nu_e + d \rightarrow p + p + e^-$ interactions produced by ^8B solar neutrinos at the Sudbury Neutrino Observatory*, *Phys. Rev. Lett.* **87** (2001) 071301 [nucl-ex/0106015] [INSPIRE].
- [10] SNO collaboration, *Direct evidence for neutrino flavor transformation from neutral current interactions in the Sudbury Neutrino Observatory*, *Phys. Rev. Lett.* **89** (2002) 011301 [nucl-ex/0204008] [INSPIRE].
- [11] DAYA BAY collaboration, *Observation of electron-antineutrino disappearance at Daya Bay*, *Phys. Rev. Lett.* **108** (2012) 171803 [arXiv:1203.1669] [INSPIRE].
- [12] DAYA BAY collaboration, *Improved measurement of electron antineutrino disappearance at Daya Bay*, *Chin. Phys. C* **37** (2013) 011001 [arXiv:1210.6327] [INSPIRE].
- [13] DOUBLE CHOOZ collaboration, *Indication of reactor $\bar{\nu}_e$ disappearance in the Double Chooz experiment*, *Phys. Rev. Lett.* **108** (2012) 131801 [arXiv:1112.6353] [INSPIRE].
- [14] RENO collaboration, *Observation of reactor electron antineutrino disappearance in the RENO experiment*, *Phys. Rev. Lett.* **108** (2012) 191802 [arXiv:1204.0626] [INSPIRE].
- [15] SUPER-KAMIOKANDE collaboration, *Search for differences in oscillation parameters for atmospheric neutrinos and antineutrinos at Super-Kamiokande*, *Phys. Rev. Lett.* **107** (2011) 241801 [arXiv:1109.1621] [INSPIRE].
- [16] M.G. Aartsen et al., *Neutrino oscillation studies with IceCube-DeepCore*, *Nucl. Phys. B* **908** (2016) 161 [INSPIRE].
- [17] G. Altarelli, N. Cabibbo, L. Maiani and R. Petronzio, *Must the new heavy lepton have its own neutrino?*, *Phys. Lett. B* **67** (1977) 463 [INSPIRE].
- [18] DONUT collaboration, *Observation of tau neutrino interactions*, *Phys. Lett. B* **504** (2001) 218 [hep-ex/0012035] [INSPIRE].
- [19] OPERA collaboration, *Discovery of τ neutrino appearance in the CNGS neutrino beam with the OPERA experiment*, *Phys. Rev. Lett.* **115** (2015) 121802 [arXiv:1507.01417] [INSPIRE].
- [20] S. Parke and M. Ross-Lonergan, *Unitarity and the three flavor neutrino mixing matrix*, *Phys. Rev. D* **93** (2016) 113009 [arXiv:1508.05095] [INSPIRE].
- [21] Z. Hu, J. Ling, J. Tang and T.C. Wang, *Global oscillation data analysis on the 3ν mixing without unitarity*, *JHEP* **01** (2021) 124 [arXiv:2008.09730] [INSPIRE].
- [22] P.B. Denton and J. Gehrlein, *New oscillation and scattering constraints on the tau row matrix elements without assuming unitarity*, *JHEP* **06** (2022) 135 [arXiv:2109.14575] [INSPIRE].

- [23] SUPER-KAMIOKANDE collaboration, *Measurement of the tau neutrino cross section in atmospheric neutrino oscillations with Super-Kamiokande*, *Phys. Rev. D* **98** (2018) 052006 [[arXiv:1711.09436](#)] [[INSPIRE](#)].
- [24] R. Mammen Abraham et al., *Tau neutrinos in the next decade: from GeV to EeV*, *J. Phys. G* **49** (2022) 110501 [[arXiv:2203.05591](#)] [[INSPIRE](#)].
- [25] KM3NeT collaboration, *Determining the neutrino mass ordering and oscillation parameters with KM3NeT/ORCA*, *Eur. Phys. J. C* **82** (2022) 26 [[arXiv:2103.09885](#)] [[INSPIRE](#)].
- [26] KM3NeT collaboration, *Letter of intent for KM3NeT 2.0*, *J. Phys. G* **43** (2016) 084001 [[arXiv:1601.07459](#)] [[INSPIRE](#)].
- [27] KM3NeT collaboration, *The KM3NeT multi-PMT optical module*, *2022 JINST* **17** P07038 [[arXiv:2203.10048](#)] [[INSPIRE](#)].
- [28] KM3NeT collaboration, *The KM3NeT data acquisition system — status and evolution*, *EPJ Web Conf.* **280** (2023) 08004 [[INSPIRE](#)].
- [29] KM3NeT collaboration, *Measurement of neutrino oscillation parameters with the first six detection units of KM3NeT/ORCA*, *JHEP* **10** (2024) 206 [[arXiv:2408.07015](#)] [[INSPIRE](#)].
- [30] ANTARES collaboration, *Monte Carlo simulations for the ANTARES underwater neutrino telescope*, *JCAP* **01** (2021) 064 [[arXiv:2010.06621](#)] [[INSPIRE](#)].
- [31] C. Andreopoulos et al., *The GENIE neutrino Monte Carlo generator*, *Nucl. Instrum. Meth. A* **614** (2010) 87 [[arXiv:0905.2517](#)] [[INSPIRE](#)].
- [32] KM3NeT collaboration, *gSeaGen: the KM3NeT GENIE-based code for neutrino telescopes*, *Comput. Phys. Commun.* **256** (2020) 107477 [[arXiv:2003.14040](#)] [[INSPIRE](#)].
- [33] M. Honda et al., *Atmospheric neutrino flux calculation using the NRLMSISE-00 atmospheric model*, *Phys. Rev. D* **92** (2015) 023004 [[arXiv:1502.03916](#)] [[INSPIRE](#)].
- [34] A.M. Dziewonski and D.L. Anderson, *Preliminary reference earth model*, *Phys. Earth Planet. Interiors* **25** (1981) 297 [[INSPIRE](#)].
- [35] J. Coelho et al., *joaoabcoelho/oscprob: v2.0.12*, *Zenodo* (2023).
- [36] G. Carminati, A. Margiotta and M. Spurio, *Atmospheric MUons from PArametric formulas: a Fast GEnerator for neutrino telescopes (MUPAGE)*, *Comput. Phys. Commun.* **179** (2008) 915 [[arXiv:0802.0562](#)] [[INSPIRE](#)].
- [37] M. de Jong and E. van Campenhout, *The probability density function of the arrival time of Čerenkov light*, [arXiv:2305.19626](#) [[INSPIRE](#)].
- [38] A. Bhattacharya et al., *Prompt atmospheric neutrino fluxes: perturbative QCD models and nuclear effects*, *JHEP* **11** (2016) 167 [[arXiv:1607.00193](#)] [[INSPIRE](#)].
- [39] ICECUBE collaboration, *Measurement of atmospheric tau neutrino appearance with IceCube DeepCore*, *Phys. Rev. D* **99** (2019) 032007 [[arXiv:1901.05366](#)] [[INSPIRE](#)].
- [40] I. Martinez-Soler and H. Minakata, *Measuring tau neutrino appearance probability via unitarity*, *Phys. Rev. D* **104** (2021) 093006 [[arXiv:2109.06933](#)] [[INSPIRE](#)].
- [41] F.J. Escrivuela et al., *On the description of nonunitary neutrino mixing*, *Phys. Rev. D* **92** (2015) 053009 [*Erratum ibid.* **93** (2016) 119905] [[arXiv:1503.08879](#)] [[INSPIRE](#)].
- [42] M. Malinsky, T. Ohlsson, Z.-Z. Xing and H. Zhang, *Non-unitary neutrino mixing and CP violation in the minimal inverse seesaw model*, *Phys. Lett. B* **679** (2009) 242 [[arXiv:0905.2889](#)] [[INSPIRE](#)].

- [43] H.-C. Han and Z.-Z. Xing, *A full parametrization of the 9×9 active-sterile flavor mixing matrix in the inverse or linear seesaw scenario of massive neutrinos*, *Nucl. Phys. B* **973** (2021) 115609 [[arXiv:2110.12705](#)] [[INSPIRE](#)].
- [44] S. Okubo, *Note on unitary symmetry in strong interaction. II Excited states of baryons*, *Prog. Theor. Phys.* **28** (1962) 24 [[INSPIRE](#)].
- [45] S. Sahoo, S. Das, A. Kumar and S.K. Agarwalla, *Constraining non-unitary neutrino mixing using matter effects in atmospheric neutrinos at INO-ICAL*, *JHEP* **09** (2024) 184 [[arXiv:2309.16942](#)] [[INSPIRE](#)].
- [46] L.S. Miranda, P. Pasquini, U. Rahaman and S. Razzaque, *Searching for non-unitary neutrino oscillations in the present T2K and NOvA data*, *Eur. Phys. J. C* **81** (2021) 444 [[arXiv:1911.09398](#)] [[INSPIRE](#)].
- [47] C.S. Fong, *Theoretical aspect of nonunitarity in neutrino oscillation*, [arXiv:2301.12960](#) [[INSPIRE](#)].
- [48] NOMAD collaboration, *Search for $\nu_\mu \rightarrow \nu_e$ oscillations in the NOMAD experiment*, *Phys. Lett. B* **570** (2003) 19 [[hep-ex/0306037](#)] [[INSPIRE](#)].
- [49] CHORUS collaboration, *Results from the CHORUS experiment at CERN*, *Nucl. Phys. B Proc. Suppl.* **85** (2000) 95 [[INSPIRE](#)].
- [50] M. Blennow et al., *Non-unitarity, sterile neutrinos, and non-standard neutrino interactions*, *JHEP* **04** (2017) 153 [[arXiv:1609.08637](#)] [[INSPIRE](#)].
- [51] ANTARES collaboration, *Measuring the atmospheric neutrino oscillation parameters and constraining the 3+1 neutrino model with ten years of ANTARES data*, *JHEP* **06** (2019) 113 [[arXiv:1812.08650](#)] [[INSPIRE](#)].
- [52] R.J. Barlow and C. Beeston, *Fitting using finite Monte Carlo samples*, *Comput. Phys. Commun.* **77** (1993) 219 [[INSPIRE](#)].
- [53] J.S. Conway, *Incorporating nuisance parameters in likelihoods for multisource spectra*, in the proceedings of the *PHYSTAT 2011*, (2011) [[DOI:10.5170/CERN-2011-006.115](#)] [[arXiv:1103.0354](#)] [[INSPIRE](#)].
- [54] F. James, *MINUIT function minimization and error analysis: reference manual version 94.1*, CERN-D-506, CERN, Geneva, Switzerland (1994) [[INSPIRE](#)].
- [55] G.D. Barr, T.K. Gaisser, S. Robbins and T. Stanev, *Uncertainties in atmospheric neutrino fluxes*, *Phys. Rev. D* **74** (2006) 094009 [[astro-ph/0611266](#)] [[INSPIRE](#)].
- [56] G.J. Feldman and R.D. Cousins, *A unified approach to the classical statistical analysis of small signals*, *Phys. Rev. D* **57** (1998) 3873 [[physics/9711021](#)] [[INSPIRE](#)].
- [57] K. Cranmer, *Practical statistics for the LHC*, in the proceedings of the *2011 European school of high-energy physics*, (2014) [[DOI:10.5170/CERN-2014-003.267](#)] [[arXiv:1503.07622](#)] [[INSPIRE](#)].
- [58] B. Efron, *Bootstrap methods: another look at the Jackknife*, *Annals Statist.* **7** (1979) 1 [[INSPIRE](#)].
- [59] S.S. Wilks, *The large-sample distribution of the likelihood ratio for testing composite hypotheses*, *Annals Math. Statist.* **9** (1938) 60 [[INSPIRE](#)].
- [60] OPERA collaboration, *Final results of the OPERA experiment on ν_τ appearance in the CNGS neutrino beam*, *Phys. Rev. Lett.* **120** (2018) 211801 [*Erratum ibid.* **121** (2018) 139901] [[arXiv:1804.04912](#)] [[INSPIRE](#)].

The KM3NeT collaboration

S. Aiello^a, A. Albert^{b,be}, A. R. Alhebsi^c, M. Alshamsi^d, S. Alves Garre^e, A. Ambrosone^{g,f}, F. Ameli^h, M. Andreⁱ, L. Aphecetche^j, M. Ardid^k, S. Ardid^k, J. Aublin^l, F. Badaracco^{n,m}, L. Bailly-Salins^o, Z. Bardačová^{q,p}, B. Baret^l, A. Bariego-Quintana^e, Y. Becherini^l, M. Bendahman^f, F. Benfenati Gualandi^{s,r}, M. Benhassi^{t,f}, M. Bennani^u, D. M. Benoit^v, E. Berbee^w, V. Bertin^d, S. Biagi^x, M. Boettcher^y, D. Bonanno^x, A. B. Bouasla^{bf}, J. Boumaaza^z, M. Bouta^d, M. Bouwhuis^w, C. Bozza^{aa,f}, R. M. Bozza^{g,f}, H. Brânzaş^{ab}, F. Bretaudeau^j, M. Breuhaus^d, R. Bruijn^{ac,w}, J. Brunner^d, R. Bruno^a, E. Buis^{ad,w}, R. Buompane^{t,f}, J. Busto^d, B. Cai^{ff}, D. Calvo^e, A. Capone^{h,ae}, F. Carenini^{s,r}, V. Carretero^{ac,w}, T. Cartraud^l, P. Castaldi^{af,r}, V. Cecchini^e, S. Celli^{h,ae}, L. Cerisy^{id,*}, M. Chabab^{ag}, A. Chen^{ah}, S. Cherubini^{ai,x}, T. Chiarusi^r, M. Circella^{aj}, R. Clark^{ak}, R. Cocimano^x, J. A. B. Coelho^l, A. Coleiro^l, A. Condorelli^{g,f}, R. Coniglione^x, P. Coyle^d, A. Creusot^l, G. Cuttone^x, R. Dallier^j, A. De Benedittis^f, G. De Wasseige^{ak}, V. Decoene^j, P. Deguire^d, I. Del Rosso^{s,r}, L. S. Di Mauro^x, I. Di Palma^{h,ae}, A. F. Díaz^{al}, D. Diego-Tortosa^x, C. Distefano^x, A. Domi^{am}, C. Donzaud^l, D. Dornic^d, E. Drakopoulou^{an}, D. Drouhin^{b,be}, J.-G. Ducoin^d, P. Duverne^l, R. Dvornický^q, T. Eberl^{am}, E. Eckerová^{q,p}, A. Eddymaoui^z, T. van Eeden^w, M. Eff^l, D. van Eijk^w, I. El Bojaddaini^{ao}, S. El Hedri^l, S. El Mentawi^d, V. Ellajosyula^{n,m}, A. Enzenhöfer^d, G. Ferrara^{ai,x}, M. D. Filipović^{ap}, F. Filippini^r, D. Franciotti^x, L. A. Fusco^{aa,f}, S. Gagliardini^{ae,h}, T. Gal^{am}, J. García Méndez^k, A. Garcia Soto^e, C. Gatus Oliver^w, N. Geißelbrecht^{id,am,*}, E. Genton^{ak}, H. Ghaddari^{ao}, L. Gialanella^{t,f}, B. K. Gibson^v, E. Giorgio^x, I. Goos^l, P. Goswami^l, S. R. Gozzini^e, R. Gracia^{am}, C. Guidi^{m,n}, B. Guillon^o, M. Gutiérrez^{aq}, C. Haack^{am}, H. van Haren^{ar}, A. Heijboer^w, L. Hennig^{am}, J. J. Hernández-Rey^e, A. Idrissi^x, W. Idrissi Ibsalhi^f, G. Illuminati^{s,r}, D. Joly^d, M. de Jong^{as,w}, P. de Jong^{ac,w}, B. J. Jung^w, P. Kalaczyński^{bg,at}, V. Kikvadze^{au}, G. Kistauri^{av,au}, C. Kopper^{am}, A. Kouchner^{aw,l}, Y. Y. Kovalev^{ax}, L. Krupa^p, V. Kueviakoe^w, V. Kulikovskiyⁿ, R. Kvatadze^{av}, M. Labalme^o, R. Lahmann^{am}, M. Lamoureux^{ak}, G. Larosa^x, C. Lastoria^{io,*}, J. Lazar^{ak}, A. Lazo^e, S. Le Stum^d, G. Lehaut^o, V. Lemaître^{ak}, E. Leonora^a, N. Lessing^e, G. Levi^{s,r}, M. Lindsey Clark^l, F. Longhitano^a, F. Magnani^d, J. Majumdar^w, L. Malerba^{n,m}, F. Mamedov^p, A. Manfreda^f, A. Manousakis^{ay}, M. Marconi^{m,n}, A. Margiotta^{s,r}, A. Marinelli^{g,f}, C. Markou^{an}, L. Martin^j, M. Mastrodicasa^{ae,h}, S. Mastroianni^f, J. Mauro^{ak}, K. C. K. Mehta^{az}, A. Meskar^{ba}, G. Miele^{g,f}, P. Migliozi^f, E. Migneco^x, M. L. Mitsou^{t,f}, C. M. Mollo^f, L. Morales-Gallegos^{t,f}, A. Moussa^{ao}, I. Mozun Mateo^o, R. Muller^r, M. R. Musone^{t,f}, M. Musumeci^x, S. Navas^{aq}, A. Nayerhoda^{aj}, C. A. Nicolau^h, B. Nkosi^{ah}, B. Ó Fearraighⁿ, V. Oliviero^{g,f}, A. Orlando^x, E. Oukacha^l, D. Paesani^x, J. Palacios González^e, G. Papalashvili^{aj,au}, V. Parisi^{m,n}, A. Parmar^o, E. J. Pastor Gomez^e, C. Pastore^{aj}, A. M. Păun^{ab}, G. E. Pāvālaš^{ab}, S. Peña Martínez^l, M. Perrin-Terrin^d, V. Pestel^o, R. Pestes^l, P. Piattelli^x, A. Plavin^{ax,bh}, C. Poirè^{aa,f}, V. Popa^{ab,1}, T. Pradier^b, J. Prado^e, S. Pulvirenti^x, C. A. Quiroz-Rangel^k, N. Randazzo^a, A. Ratnani^{bb}, S. Razzaque^{bc}, I. C. Rea^f, D. Real^e, G. Riccobene^x, A. Romanov^{m,n,o}, E. Ros^{ax}, A. Šaina^e, F. Salesa Greus^e, D. F. E. Samtleben^{as,w}, A. Sánchez Losa^e, S. Sanfilippo^x, M. Sanguineti^{m,n}, D. Santonocito^x, P. Sapienza^x, M. Scarnera^{ak,l}, J. Schnabel^{am}, J. Schumann^{am}, H. M. Schutte^y, J. Seneca^w, N. Sennan^{ao}, P. Sevlé^{ak}, I. Sgura^{aj}, R. Shanidze^{au}, A. Sharma^l, Y. Shitov^p, F. Šimković^q, A. Simonelli^f, A. Sinopoulou^a, B. Spisso^f, M. Spurio^{s,r}, D. Stavropoulos^{an}, I. Štekl^p, M. Taiuti^{m,n}, G. Takadze^{au}, Y. Tayalati^{z,bb}, H. Thiersen^y, S. Thoudam^c, I. Tosta e Melo^{a,ai}, B. Trocmé^l, V. Tsourapis^{an}, A. Tudorache^{h,ae}, E. Tzamariudaki^{an}, A. Ukleja^{ba}, A. Vacheret^o, V. Valsecchi^x, V. Van Elewyck^{aw,l}, G. Vannoye^{d,n}, G. Vasileiadis^{bd},

F. Vazquez de Sola^w, A. Veutro^{h,ae}, S. Viola^x, D. Vivolo^{t,f}, A. van Vliet^c, E. de Wolf^{ac,w},
I. Lhenry-Yvon^l, S. Zavatarelliⁿ, A. Zegarelli^{h,ae}, D. Zito^x, J. D. Zornoza^e, J. Zúñiga^e, N. Zywucka^y

^a INFN, Sezione di Catania, (INFN-CT) Via Santa Sofia 64, Catania, 95123 Italy

^b Université de Strasbourg, CNRS, IPHC UMR 7178, F-67000 Strasbourg, France

^c Khalifa University of Science and Technology, Department of Physics, PO Box 127788, Abu Dhabi, United Arab Emirates

^d Aix Marseille Univ, CNRS/IN2P3, CPPM, Marseille, France

^e IFIC — Instituto de Física Corpuscular (CSIC — Universitat de València), c/Catedrático José Beltrán, 2, 46980 Paterna, Valencia, Spain

^f INFN, Sezione di Napoli, Complesso Universitario di Monte S. Angelo, Via Cintia ed. G, Napoli, 80126 Italy

^g Università di Napoli “Federico II”, Dip. Scienze Fisiche “E. Pancini”, Complesso Universitario di Monte S. Angelo, Via Cintia ed. G, Napoli, 80126 Italy

^h INFN, Sezione di Roma, Piazzale Aldo Moro 2, Roma, 00185 Italy

ⁱ Universitat Politècnica de Catalunya, Laboratori d’Aplicacions Bioacústiques, Centre Tecnològic de Vilanova i la Geltrú, Avda. Rambla Exposició, s/n, Vilanova i la Geltrú, 08800 Spain

^j Subatech, IMT Atlantique, IN2P3-CNRS, Nantes Université, 4 rue Alfred Kastler — La Chantrerie, Nantes, BP 20722 44307 France

^k Universitat Politècnica de València, Instituto de Investigación para la Gestión Integrada de las Zonas Costeras, C/ Paranimf, 1, Gandia, 46730 Spain

^l Université Paris Cité, CNRS, Astroparticule et Cosmologie, F-75013 Paris, France

^m Università di Genova, Via Dodecaneso 33, Genova, 16146 Italy

ⁿ INFN, Sezione di Genova, Via Dodecaneso 33, Genova, 16146 Italy

^o LPC CAEN, Normandie Univ, ENSICAEN, UNICAEN, CNRS/IN2P3, 6 boulevard Maréchal Juin, Caen, 14050 France

^p Czech Technical University in Prague, Institute of Experimental and Applied Physics, Husova 240/5, Prague, 110 00 Czech Republic

^q Comenius University in Bratislava, Department of Nuclear Physics and Biophysics, Mlynska dolina F1, Bratislava, 842 48 Slovak Republic

^r INFN, Sezione di Bologna, v.le C. Berti-Pichat, 6/2, Bologna, 40127 Italy

^s Università di Bologna, Dipartimento di Fisica e Astronomia, v.le C. Berti-Pichat, 6/2, Bologna, 40127 Italy

^t Università degli Studi della Campania “Luigi Vanvitelli”, Dipartimento di Matematica e Fisica, viale Lincoln 5, Caserta, 81100 Italy

^u LPC, Campus des Cézeaux 24, avenue des Landais BP 80026, Aubière Cedex, 63171 France

^v E. A. Milne Centre for Astrophysics, University of Hull, Hull, HU6 7RX, United Kingdom

^w Nikhef, National Institute for Subatomic Physics, PO Box 41882, Amsterdam, 1009 DB Netherlands

^x INFN, Laboratori Nazionali del Sud, (LNS) Via S. Sofia 62, Catania, 95123 Italy

^y North-West University, Centre for Space Research, Private Bag X6001, Potchefstroom, 2520 South Africa

^z University Mohammed V in Rabat, Faculty of Sciences, 4 av. Ibn Battouta, B.P. 1014, R.P. 10000 Rabat, Morocco

^{aa} Università di Salerno e INFN Gruppo Collegato di Salerno, Dipartimento di Fisica, Via Giovanni Paolo II 132, Fisciano, 84084 Italy

^{ab} ISS, Atomistilor 409, Măgurele, RO-077125 Romania

^{ac} University of Amsterdam, Institute of Physics/IHEF, PO Box 94216, Amsterdam, 1090 GE Netherlands

^{ad} TNO, Technical Sciences, PO Box 155, Delft, 2600 AD Netherlands

^{ae} Università La Sapienza, Dipartimento di Fisica, Piazzale Aldo Moro 2, Roma, 00185 Italy

^{af} Università di Bologna, Dipartimento di Ingegneria dell’Energia Elettrica e dell’Informazione “Guglielmo Marconi”, Via dell’Università 50, Cesena, 47521 Italia

^{ag} Cadi Ayyad University, Physics Department, Faculty of Science Semlalia, Av. My Abdellah, P.O.B. 2390, Marrakech, 40000 Morocco

^{ah} University of the Witwatersrand, School of Physics, Private Bag 3, Johannesburg, Wits 2050 South Africa

- ^{ai} *Università di Catania, Dipartimento di Fisica e Astronomia “Ettore Majorana”, (INFN-CT) Via Santa Sofia 64, Catania, 95123 Italy*
- ^{aj} *INFN, Sezione di Bari, via Orabona, 4, Bari, 70125 Italy*
- ^{ak} *UCLouvain, Centre for Cosmology, Particle Physics and Phenomenology, Chemin du Cyclotron, 2, Louvain-la-Neuve, 1348 Belgium*
- ^{al} *University of Granada, Department of Computer Engineering, Automation and Robotics / CITIC, 18071 Granada, Spain*
- ^{am} *Friedrich-Alexander-Universität Erlangen-Nürnberg (FAU), Erlangen Centre for Astroparticle Physics, Nikolaus-Fiebiger-Straße 2, 91058 Erlangen, Germany*
- ^{an} *NCSR Demokritos, Institute of Nuclear and Particle Physics, Ag. Paraskevi Attikis, Athens, 15310 Greece*
- ^{ao} *University Mohammed I, Faculty of Sciences, BV Mohammed VI, B.P. 717, R.P. 60000 Oujda, Morocco*
- ^{ap} *Western Sydney University, School of Computing, Engineering and Mathematics, Locked Bag 1797, Penrith, NSW 2751 Australia*
- ^{aq} *University of Granada, Dpto. de Física Teórica y del Cosmos & C.A.F.P.E., 18071 Granada, Spain*
- ^{ar} *NIOZ (Royal Netherlands Institute for Sea Research), PO Box 59, Den Burg, Texel, 1790 AB, the Netherlands*
- ^{as} *Leiden University, Leiden Institute of Physics, PO Box 9504, Leiden, 2300 RA Netherlands*
- ^{at} *AGH University of Krakow, Center of Excellence in Artificial Intelligence, Al. Mickiewicza 30, Krakow, 30-059 Poland*
- ^{au} *Tbilisi State University, Department of Physics, 3, Chavchavadze Ave., Tbilisi, 0179 Georgia*
- ^{av} *The University of Georgia, Institute of Physics, Kostava str. 77, Tbilisi, 0171 Georgia*
- ^{aw} *Institut Universitaire de France, 1 rue Descartes, Paris, 75005 France*
- ^{ax} *Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, 53121 Bonn, Germany*
- ^{ay} *University of Sharjah, Sharjah Academy for Astronomy, Space Sciences, and Technology, University Campus — POB 27272, Sharjah, — United Arab Emirates*
- ^{az} *AGH University of Krakow, Faculty of Physics and Applied Computer Science, Reymonta 19, Krakow, 30-059 Poland*
- ^{ba} *National Centre for Nuclear Research, 02-093 Warsaw, Poland*
- ^{bb} *School of Applied and Engineering Physics, Mohammed VI Polytechnic University, Ben Guerir, 43150, Morocco*
- ^{bc} *University of Johannesburg, Department Physics, PO Box 524, Auckland Park, 2006 South Africa*
- ^{bd} *Laboratoire Univers et Particules de Montpellier, Place Eugène Bataillon — CC 72, Montpellier Cédex 05, 34095 France*
- ^{be} *Université de Haute Alsace, rue des Frères Lumière, 68093 Mulhouse Cedex, France*
- ^{bf} *Université Badji Mokhtar, Département de Physique, Faculté des Sciences, Laboratoire de Physique des Rayonnements, B. P. 12, Annaba, 23000 Algeria*
- ^{bg} *AstroCeNT, Nicolaus Copernicus Astronomical Center, Polish Academy of Sciences, Rektorska 4, Warsaw, 00-614 Poland*
- ^{bh} *Harvard University, Black Hole Initiative, 20 Garden Street, Cambridge, MA 02138 U.S.A.*

* Corresponding author

¹ Deceased.